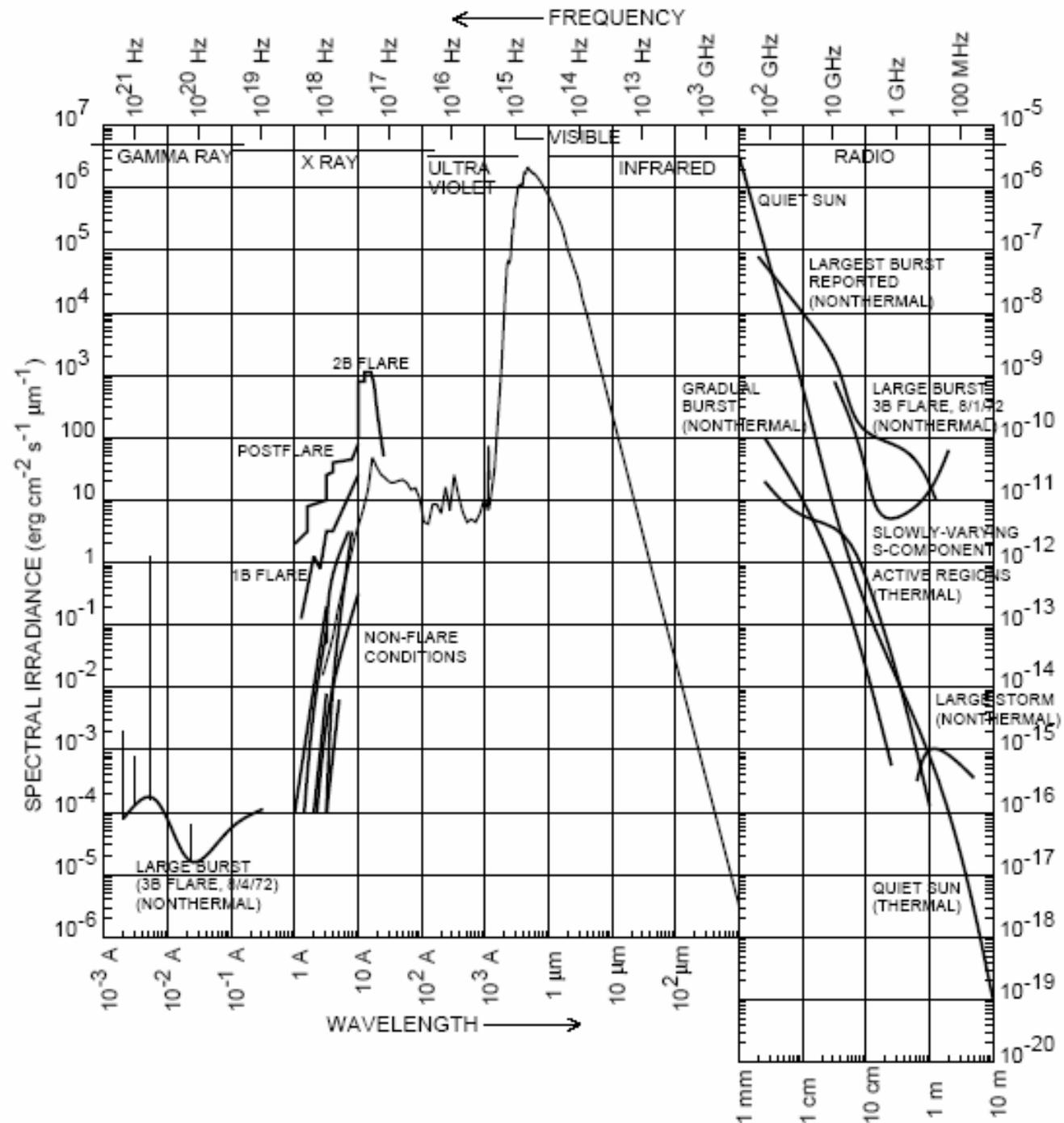
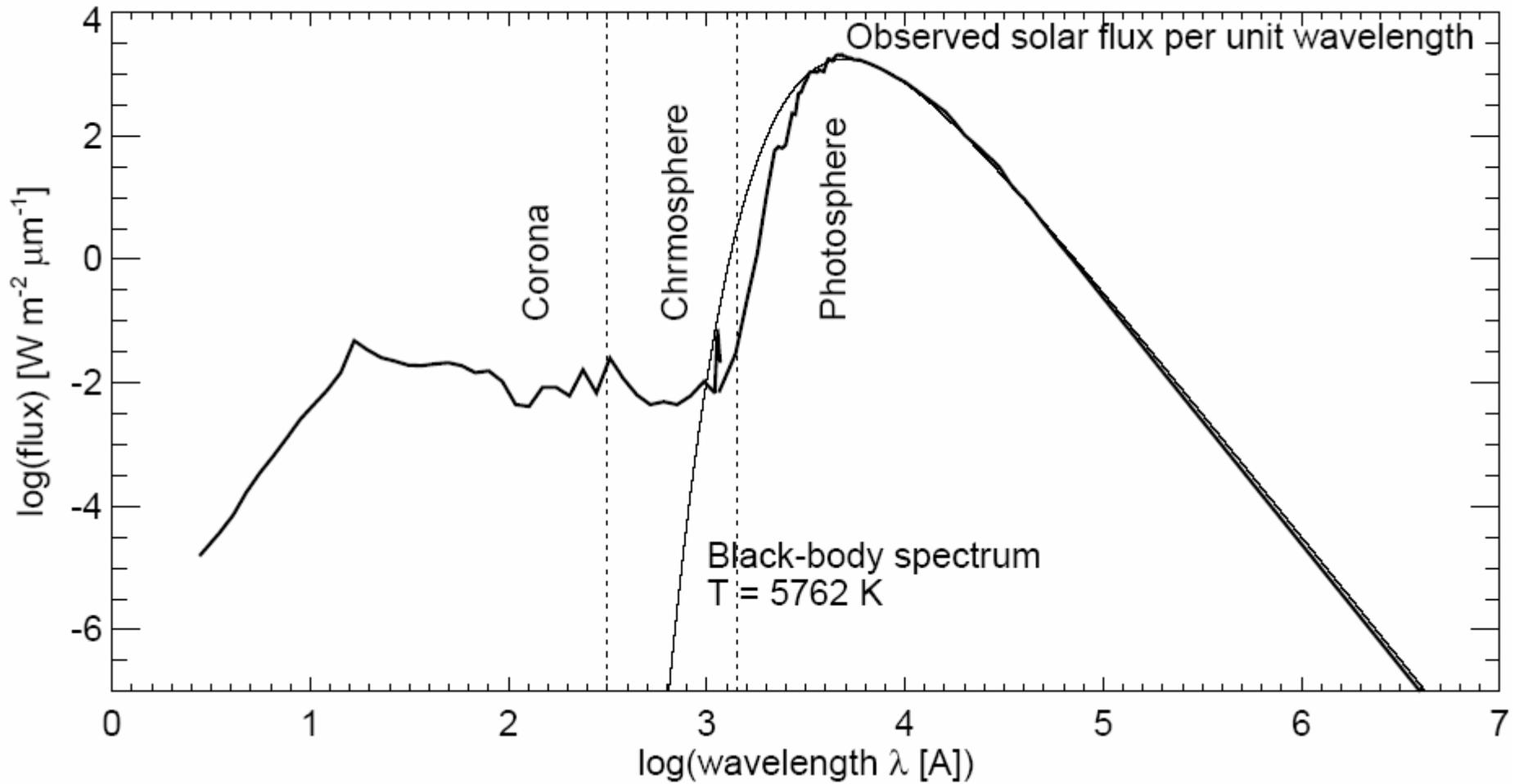


¿De qué color es el Sol?



$$B_{\lambda}(T) = \frac{2hc^2n_{\nu}^2}{\lambda^5} \frac{1}{[\exp (hc/\lambda k_B T) - 1]}$$



Ventanas espectrales para identificar estructuras con observaciones remotas. Ingredientes.

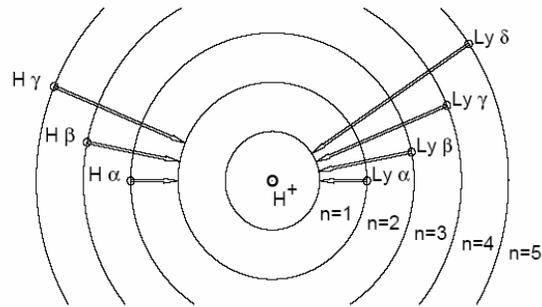


Figure 2.7: Bohr–Sommerfeld model of the hydrogen atom. The lowest quantum mechanical orbits $n = 1, 2, 3, 4, 5, \dots$, are indicated, along with the transitions of the shortest wavelengths, the Lyman series (Ly α , Ly β , ...) and the Balmer series (H α , H β , ...), that are emitted when an electron makes the indicated transitions.

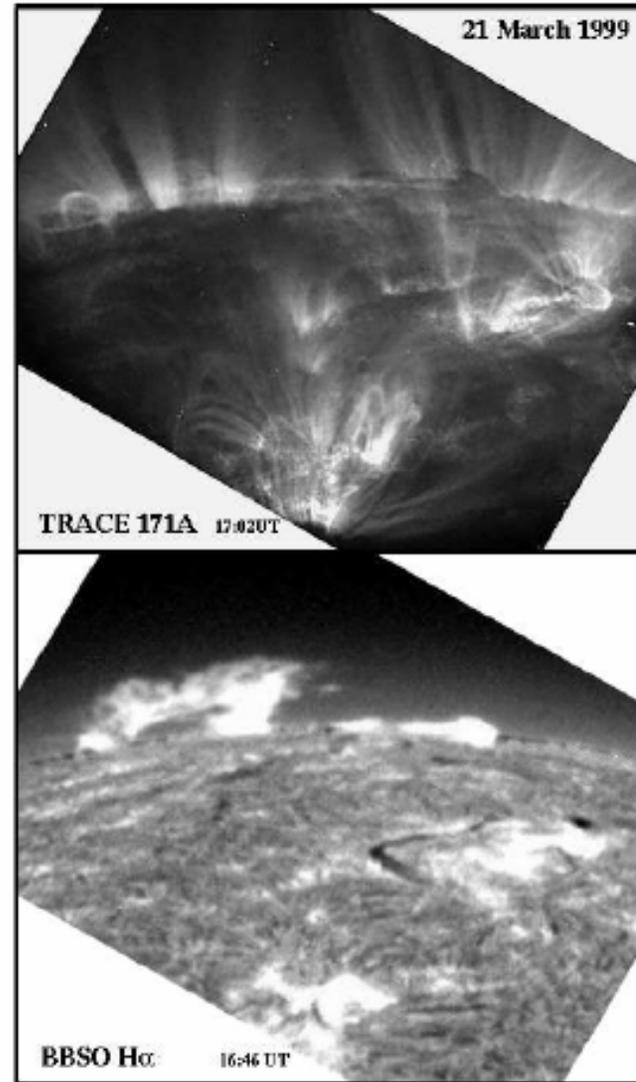
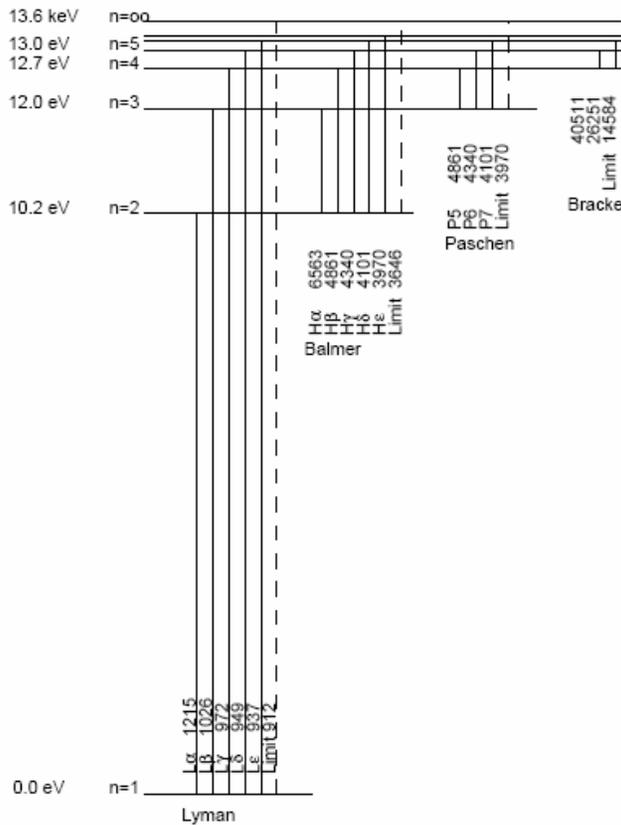


Figure 1.10: These images (taken on 1999 March 21) compare the corona seen in EUV (top panel: *TRACE*, 171 Å, $T=1$ MK) and the chromosphere seen in H α (bottom panel: *Big Bear Solar Observatory (BBSO)*, $T=10,000$ K). The cool filaments (on the disk) and prominences (above the limb) show up as bright structures in H α (bottom frame), but as dark, absorbing features in EUV (top frame) (courtesy of *TRACE* and *BBSO*).